

Time-dependent compactification to de Sitter space: a no-go theorem

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ABSTRACT: It is known that the Einstein gravitational field equations in $D > 4$ spacetime dimensions have no time-independent non-singular compactification solutions to de Sitter space if the D -dimensional stress tensor satisfies the Strong Energy Condition (SEC). Here we show, by example, that the SEC alone does not exclude time-dependent non-singular compactifications to de Sitter space, in Einstein conformal frame. However, this possibility is excluded by the combined SEC and Dominant Energy Condition (DEC) because the DEC forces a time-evolution towards a singular D -metric.

KEYWORDS: Classical Theories of Gravity, Spacetime Singularities, Supergravity Models, Energy conditions

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1 Introduction

The observed accelerated expansion of the Universe is consistent with a de Sitter (dS) phase, on the assumption of approximate homogeneity and isotropy. This fact is in tension with string/M-theory because the low energy effective supergravity theories in 10/11 dimensions do not admit time-independent compactifications to de Sitter space. In most cases this follows from a simple no-go theorem (originally due to Gibbons [1, 2] and rediscovered in a String/M-theory context by Maldacena and Nuñez [3]) that rules out this possibility for a stress tensor satisfying the Strong Energy Condition (SEC).

Some types of singularity are innocuous in String Theory and it may be that allowance for them also allows compactifications to de Sitter space (which we abbreviate to “dS compactifications”). Proposals along these lines have been made, notably [4], but no consensus on their validity has yet been reached, and there is some recent contrary evidence that consigns dS compactifications to a cosmological “swampland” [5].

Another way that the Gibbons-Maldacena-Nuñez (GMN) no-go theorem might be circumvented is to allow for time-dependence of the compact space metric. In this case there is an ambiguity in the metric on the lower-dimensional spacetime, with different choices related by a field redefinition involving time-dependent scalar fields. This ambiguity is resolved by a choice of “conformal frame”, and the “Einstein frame” (which results in the absence of any time-dependent function of scalar fields multiplying the Ricci scalar in the lower-dimensional Einstein-Hilbert integrand) is the standard choice.

As we emphasised in a previous work [6], the Einstein-frame condition involves an integration over the compact space but implementation of it has always (to our knowledge) involved a restriction on the integrand that is sufficient but not necessary. If the Einstein-frame condition is implemented in this “unaveraged” way then it is possible to prove that the SEC rules out even time-dependent (non-singular) dS compactifications but this result does not apply more generally, as we show here by means of an explicit 5D example with a compact space that is topologically a circle. The Einstein-frame condition is satisfied in this example, but *not* in its “unaveraged” form.

This counter-example to the would-be theorem that non-singular dS compactifications are forbidden by the higher-dimensional SEC is, however, non-physical: the 5D stress tensor does not satisfy the Dominant Energy condition (DEC), which is required by causality. This suggests that time-dependent non-singular dS compactifications may be ruled out by the SEC and DEC combined. In fact, our main result is that the DEC alone is sufficient for this purpose if the compact space metric is strictly time-dependent, because it then implies a singularity of the higher-dimensional metric.

If this new no-go theorem is combined with the GMN no-go theorem then we almost have a proof that non-singular dS compactifications are excluded by the SEC and DEC combined. However, the general ansatz for a time-dependent dS compactification involves not only a metric on the compact space but also a “warp factor”, and it may be that this is time-dependent even if the compact space metric is not. In this case, which has never previously been considered, we show that the DEC implies an evolution towards a discontinuous warp factor. The corresponding singular D -metric might be reached only asymptotically, but we still expect conditions required for the low-energy validity of Einstein’s field equations to be violated at some finite time.

2 Warped cosmological compactifications

It is useful to consider the issues involved in the larger context of time-dependent compactifications from D dimensions to a general homogeneous and isotropic (FLRW) spacetime of dimension $d < D$, for which the standard form of the metric is

$$ds_{\text{FLRW}}^2 \equiv g_{\mu\nu} dx^\mu dx^\nu = -dt^2 + S^2(t) \bar{g}_{ij} dx^i dx^j, \quad (2.1)$$

where $S(t)$ is the scale factor as a function of FLRW time, and \bar{g}_{ij} is the metric in local coordinates $\{x^i; i = 1, \dots, d-1\}$ for a maximally-symmetric $(d-1)$ -space with constant curvature k . There is a time-slicing of the dS universe for which the metric takes this FLRW form for any value of k . For $k = 0$, for example, we have

$$S = e^{Ht} \quad (k = 0) \quad (2.2)$$

for constant H . The FLRW metric is a solution of the d -dimensional Einstein equations for a perfect-fluid source, specified by an energy density and pressure. For a linear equation of state, the pressure to energy density ratio is a constant w , and in this context the SEC (in d dimensions) is equivalent to $w \leq -1 + 2/(d-1)$ and the DEC (in d dimensions) is equivalent to $|w| \leq 1$. The $w = -1$ case corresponds to the dS universe. The $w > -1$ cases correspond to FLRW universes with $S \sim t^\eta$, for some constant η , and $\eta \leq 1$ if the (d -dimensional) SEC is satisfied.

Our starting point will be a D -dimensional manifold that is topologically a product of this FLRW spacetime with a compact n -dimensional manifold B (so $D = d + n$) and, following [6], we consider a general D -metric of the form

$$ds_D^2 = \Omega^2(y; t) ds_{\text{FLRW}}^2 + h_{\alpha\beta}(y; t) dy^\alpha dy^\beta, \quad (2.3)$$

where $h_{\alpha\beta}$ is the metric on B in local coordinates $\{y^\alpha; \alpha = 1, \dots, n\}$, and Ω is a nowhere-zero ‘warp factor’: a scalar function on B . The condition for ds_{FLRW}^2 to be an Einstein frame metric for the effective d -dimensional gravity theory is [6]

$$\int_B d^n y \sqrt{\det h} \Omega^{d-2} = G_D/G_d, \quad (2.4)$$

where the constant on the right-hand side is the ratio of the Newton constants in the higher and lower spacetime dimensions.

Taking a time derivative of the Einstein-frame condition we deduce that

$$0 = \int_B d^n y \sqrt{\det h} \Omega^{d-2} X \equiv \langle X \rangle \quad (2.5)$$

where

$$X = \frac{1}{2} \text{tr} \left(h^{-1} \dot{h} \right) + (d-2) \left(\dot{\Omega}/\Omega \right). \quad (2.6)$$

Following [6], we shall refer to $\langle X \rangle = 0$ as the “first-order Einstein-frame condition”. Obviously, $X \equiv 0$ is not required for $\langle X \rangle = 0$, but it does suffice.

We are interested in D -metrics of the above type that satisfy the Einstein field equations; for an appropriate choice of units these equations are

$$G_{MN} = T_{MN} \quad (M, N = 0, 1, \dots, D-1). \quad (2.7)$$

We do not need to solve these equations. Instead, we compute the Einstein tensor on the left hand side directly from the metric of (2.3). We then identify this with the stress tensor, which is therefore expressed in terms of the scale factor $S(t)$, the warp factor $\Omega(t, \mathbf{y})$ and the compact-space metric $h_{\alpha\beta}(t, \mathbf{y})$. By construction, this stress tensor supports the given D -dimensional spacetime; for a specific D -metric it is a specific function of time and the coordinates of B . One could ask what specific form of matter has this stress tensor, but an answer to this question is not required for a determination of whether appropriate energy conditions are satisfied.

We first focus on the SEC. This is a condition on the stress tensor that, given the Einstein equations (2.7), is equivalent to positivity of the time-time component of the D -dimensional Ricci tensor R_{MN} . A computation yields

$$\begin{aligned} R_{00} = & -(d-1) \left[(\ddot{S}/S) + (\ddot{\Omega}/\Omega) - (\dot{\Omega}/\Omega)^2 + (\dot{\Omega}/\Omega)(\dot{S}/S) \right] - \frac{1}{2} \text{tr} \left(h^{-1} \ddot{h} \right) \\ & + \frac{1}{2} (\dot{\Omega}/\Omega) \text{tr} \left(h^{-1} \dot{h} \right) + \frac{1}{4} \text{tr} \left(h^{-1} \dot{h} \right)^2 + (d-1) |\nabla \Omega|^2 + \Omega \nabla^2 \Omega, \end{aligned} \quad (2.8)$$

where ∇ represents the covariant derivative in the internal B space with metric $h_{\alpha\beta}$. It is implicit here that $\nabla \Omega$ and $\nabla^2 \Omega$ are defined. More, generally, we will assume that all components of the D -metric are continuous and at least twice differentiable, with respect to both time and local coordinates for the compact space B . Also, as is customarily understood in this context, we assume that B has no boundary.

Consider the case of a time-independent metric on B ; i.e. $\dot{h}_{\alpha\beta} \equiv 0$. In principle the warp factor Ω^2 could still be time-dependent, and we shall investigate this possibility later,

but it is excluded if the first-order Einstein-frame condition (2.5) is implemented in an “unaveraged” way by setting $X \equiv 0$. It is also excluded by the restriction to “time-independent compactifications”. In this case, $R_{00} \geq 0$ is equivalent to

$$d(d-1) \left(\sqrt{\det h} \Omega^{d-2} \right) \ddot{S}/S \leq \sqrt{\det h} \nabla^2 \Omega^d, \quad (2.9)$$

and integration over B yields $\ddot{S} \leq 0$, which excludes accelerated expansion, and hence a dS universe. This is essentially the GMN theorem.

3 A time-dependent de Sitter compactification

We shall now show, by example, that the D -dimensional SEC does not forbid *time-dependent* non-singular dS compactifications. We start from a 5D metric of the form

$$ds_5^2 = \Omega^2(t+y) ds_{dS}^2 + \varphi^2(t+y) dy^2 \quad (y \sim y + 2\pi L). \quad (3.1)$$

The compact space has the topology of a circle, and Ω and φ are non-zero periodic functions of $t+y$. A calculation, using the flat-slicing of dS with $k=0$, shows that

$$\begin{aligned} R_{00} = & -3H^2 - (\ddot{\varphi}/\varphi) + (\dot{\varphi}/\varphi)(\dot{\Omega}/\Omega) - 3(\ddot{\Omega}/\Omega) + 3(\dot{\Omega}/\Omega)^2 + H(\dot{\Omega}/\Omega) \\ & + (\Omega/\varphi)^2 \left[(\ddot{\Omega}/\Omega) - (\dot{\varphi}/\varphi)(\dot{\Omega}/\Omega) + (\dot{\Omega}/\Omega)^2 \right]. \end{aligned} \quad (3.2)$$

The Einstein-frame condition is

$$\int_0^{2\pi L} dy \varphi(t+y) \Omega^2(t+y) = G_5/G_4, \quad (3.3)$$

but this fixes only the scale of the product function $\varphi\Omega^2$ because the left hand side is time-independent as a consequence of periodicity. As explained earlier the Einstein-frame condition implies its “first-order” variant $\langle X \rangle = 0$, even though $X \not\equiv 0$ in this example.

Let us choose

$$\Omega = 2A \{1 + a \sin[(t+y)/L]\}, \quad \varphi = A \{1 + 2a \sin[(t+y)/L]\}, \quad (3.4)$$

where the dimensionless constant A is chosen to satisfy (3.3). The dimensionless constant a must satisfy $2|a| < 1$ to ensure that Ω and φ have no zeros. For $a=0$ we have $R_{00} = -3H^2 < 0$, so the SEC is violated. Otherwise, R_{00} is periodic in $t+y$ and hence an oscillating function of y for given t . The SEC will be satisfied if this function is positive over one period of its argument and it is possible to choose a and LH such that this is the case. An example is $a=1/5$ and $LH=0.1$, for which R_{00} is plotted in figure 1, in units for which $L=1$; we find that R_{00} remains positive for $0 < LH < 0.24$. Similar examples can be found for toroidal compactification from any dimension $D > 4$.

To summarise: the SEC does not forbid time-dependent compactification to de Sitter space. However, as will be shown below, the 5D stress tensor that supports this SEC-compliant solution of the 5D Einstein equations does not satisfy the DEC.

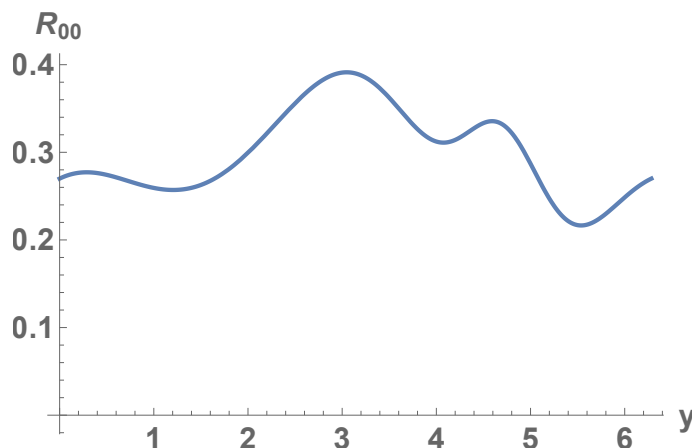


Figure 1. R_{00} of (3.2) as periodic function of y with period 2π ; computed for $L = 1$ from (3.4) with $a = 0.2$ and $H = 0.1$.

4 The DEC and a no-go theorem

According to (2.7), the $d \times d$ block of the Einstein tensor G_{MN} can be identified with the $T_{\mu\nu}$ components ($\mu, \nu = 0, 1, \dots, d-1$) of the D -dimensional stress tensor. The result takes the ideal fluid form expected from the FLRW isometries:

$$T_{00} = -\rho g_{00}, \quad T_{ij} = P g_{ij}, \quad i, j = 1, \dots, d-1. \quad (4.1)$$

It should be appreciated here that the density ρ and pressure P are functions on the D -dimensional spacetime (in contrast to the energy density and pressure in the d -dimensional FLRW spacetime, which are functions only of time). The DEC requires ρ to be greater than, or equal to, the absolute value of every other component of the (D -dimensional) stress tensor. In particular, it requires $\rho \geq |P|$, which is equivalent to $\rho \pm P \geq 0$ for both choices of sign. These combinations are

$$\begin{aligned} \rho - P &= (d-2)(\ddot{S}/S) + (d-2)^2(\dot{S}/S)^2 + (2d-3)(\dot{S}/S)X + \dot{X} + X^2 \\ &\quad - (d-1) [2\Omega \nabla^2 \Omega + (d-2)|\nabla \Omega|^2] + R(h) + (d-2)^2(k/S^2), \\ \rho + P &= -(d-2)(\ddot{S}/S) + (d-2)(\dot{S}/S)^2 + (\dot{S}/S)X - \dot{X} + \frac{1}{(d-2)}X^2 \\ &\quad - \frac{1}{4(d-2)} \left[\text{tr} \left(h^{-1} \dot{h} \right) \right]^2 - \frac{1}{4} \text{tr} \left(h^{-1} \dot{h} \right)^2 + (d-2) \frac{k}{S^2}, \end{aligned} \quad (4.2)$$

where $R(h)$ is the Ricci scalar for the compact space metric, and X was defined in (2.6). Although X has zero average over B (with weighting by $\sqrt{\det h} \Omega^{d-2}$) this is not generally true of \dot{X} since

$$\langle \dot{X} \rangle = -\langle X^2 \rangle. \quad (4.3)$$

Notice that $\rho + P$ has no term involving either $\nabla\Omega$ or $R(h)$! In the dS case, for the flat-slicing choice (2.2), the DEC inequality $\rho + P \geq 0$ is¹

$$-\dot{X} + HX + \frac{1}{(d-2)}X^2 \geq \frac{1}{4(d-2)}\left[\text{tr}\left(h^{-1}\dot{h}\right)\right]^2 + \frac{1}{4}\text{tr}\left[\left(h^{-1}\dot{h}\right)^2\right]. \quad (4.4)$$

As we are concerned with expanding universes, we may assume here that $H > 0$. The left-hand side is zero if $X \equiv 0$, which is a consequence of an “unaveraged” implementation of the Einstein-frame condition; in this case the inequality can only be satisfied if the compact space metric is time-independent. But then we have a time-independent dS compactification, which violates the SEC.

The fact that $\langle X \rangle = 0$ tells us that X will take both positive and negative values on B when $X \not\equiv 0$. Let us consider the implications of this for our dS compactification of the previous section, for which X and \dot{X} are non-zero periodic functions. In one period X will have at least two zeros (exactly two in our example) and if \dot{X} is negative at one zero then it must be positive at some other zero of X . It follows that there is a point on B (a circle in our example) for which $X = 0$ and $\dot{X} > 0$; but the left-hand side of (4.4) is then negative, which violates the inequality. The stress tensor implicit in our example therefore violates the DEC.

In general, at any given time t_0 , there will be a region of B in which $X > 0$, call it B_+ , and a region in which $X < 0$, call it B_- . There will also be points, and possibly a region, where $X = 0$; call this set B_0 . This gives us a partition of B into a union of disjoint sets:

$$B = B_- \cup B_0 \cup B_+, \quad (4.5)$$

where B_{\pm} are empty sets only when $X \equiv 0$. This partition will, in general, be time-dependent.

Let us now suppose that the compact space metric is strictly time-dependent. In this case the DEC inequality implies the strict inequality

$$-(d-2)\dot{X} + (d-2)HX + X^2 > 0. \quad (4.6)$$

Of course, it could happen that the time-dependent compact space metric evolves to one that is time-independent, such that this strict inequality is replaced by one that allows equality, but we will deal with this possibility below. For now we assume that the strict inequality holds after any sufficiently large time T . It then follows that $\dot{X} < 0$ on B_0 for $t > T$. This implies that all points in B_0 at any given time $t = t_0 > T$ will be in B_- at $t = t_0 + dt$. However, by continuity, we will also have $\dot{X} < 0$ in some neighbourhood of every boundary point of B_0 . Points in such a neighbourhood that are not in B_0 will be in B_- or B_+ , and those in B_- will remain in B_- , but those in B_+ that are sufficiently close to the boundary with B_0 will move into $B_0 \cup B_-$. Since this is true for all $t_0 > T$, there will be a flow of points from B_+ to $B_0 \cup B_-$, leading to a monotonic decrease in the volume of B_+ . Recalling that the average of X on B is zero, we see that a delta-function

¹The $k \neq 0$ time-slicings of dS, which require a different scale function $S(t)$, yield the same result but with H replaced by either $H \tanh Ht$ or $H \coth Ht$, which has no effect on the arguments to follow.

type singularity of X must form if the volume of B_+ shrinks to zero, which is inevitable unless $\dot{X} \rightarrow 0$ on B_0 , and this can happen only if $\dot{h} \rightarrow 0$ too.

If $\dot{h} \rightarrow 0$ then we have two cases to consider: either the D -metric evolves to one for which $X \equiv 0$ (which requires $\dot{\Omega} \rightarrow 0$) or it evolves to one for which $X \neq 0$. The former possibility presumes a time-evolution towards a time-independent dS compactification, which is ruled out by the SEC. This conclusion is immediate if the time-independent dS compactification is reached at finite time, but continuity implies that it is still true if we suppose that it is reached only asymptotically as $t \rightarrow \infty$.

This leaves only the possibility that the D -metric is, or evolves to, one for which the compact space metric is time-independent but the warp factor is still time-dependent, in which case $X = (d-2)\dot{\Omega}/\Omega$. The DEC inequality (4.4) is now

$$-(d-2)\dot{X} + (d-2)HX + X^2 \geq 0. \quad (4.7)$$

On B_0 this implies that $\dot{X} \leq 0$. Unless this inequality is saturated, there will again be a monotonic decrease in the volume of B_+ but now we may suppose, without violating the inequality, that the D -metric evolves to one for which B_+ has some non-zero volume and $\dot{X} = 0$ on B_0 , from which time onwards the partition (4.5) is time-independent.

In the context of this time-independent partition of B , consider points in B_- in the vicinity of B_0 , where $X = -\epsilon$ for small positive ϵ . In this case (4.7) becomes

$$-\dot{X} - H\epsilon + O(\epsilon^2) \geq 0, \quad (4.8)$$

which implies that $\dot{X} < 0$ in this region for $t > t_0$. This implies in turn that the distance between the surface $X = -\epsilon$ and the B_0/B_- boundary, where $X = 0$, must shrink to zero, at least asymptotically, leading to a discontinuity of X . This discontinuity must either be reached in finite time, implying a singular D -metric, or asymptotically as $t \rightarrow \infty$. In the latter case the D -metric itself is not singular, but its singular limit implies a breakdown of the approximation implicit in our use of the D -dimensional Einstein field equations to discuss cosmological compactifications.

5 Summary and discussion

Gravity is an attractive force only for matter satisfying the SEC, but this is not a fundamental physical requirement. A more important condition physically, for various reasons, is the DEC. The dark energy generally presumed to be the cause of the observed dS-like accelerated expansion of the universe is the form of “matter” that maximally violates the SEC while still satisfying the DEC. However, effective supergravity theories for String/M-Theory in dimensions $D = 10, 11$ have stress tensors that satisfy the SEC, and this presents obstacles to the idea that a 4-dimensional dS universe could arise from compactification. One such obstacle is the GMN no-go theorem for time-independent dS compactifications.

One might suppose, for any non-singular cosmological compactification, that the SEC is necessarily satisfied in the lower dimension if it is satisfied in the higher dimension. If this were true it would imply that compactification to a generic FLRW spacetime is possible only

if its expansion is non-accelerating. However, transient acceleration is typical, for reasons reviewed in [7], where it was conjectured that late-time acceleration might still be excluded. This conjecture was proved by Teo for vacuum solutions² of the higher-dimensional Einstein equations [8], but we recently exhibited counterexamples for a non-zero stress tensor satisfying both the SEC and the DEC in the higher dimension [6]. Specifically, we considered compactifications to power-law FLRW universes, for which the expansion is driven by perfect fluid matter with a linear equation of state, and we showed that the higher-dimensional SEC imposes a lower bound on the pressure to energy density ratio w ; for a 4-dimensional universe this bound is $w > -1/2$, so the expansion is accelerating for $w \in (-1/2, -1/3)$.

The SEC bound $w > -1/2$ was derived in [6] in the context of generic time-dependent compactifications to FLRW universes in Einstein conformal frame, but with this Einstein-frame condition imposed in the “unaveraged” way that has hitherto been standard. However, this implementation of the Einstein-frame condition is unnecessarily restrictive, and we have explored here some additional possibilities that arise when this restriction is relaxed. An obvious question is whether there are additional possibilities with $w < -1/2$. We have not attempted to fully answer this question but we have shown that even dS compactifications, for which $w = -1$, are not excluded by the SEC, although the DEC was violated in our example.

Thus, the SEC is a remarkably weak condition on cosmological compactifications once time-dependence is allowed. Equally remarkable is the power of the DEC in this context: even though it is never violated in the lower dimension, it forces the higher-dimensional metric to evolve towards a singularity. If we demand non-singularity then the combined SEC and DEC in the higher-dimension prevent compactification to dS irrespective of whether this compactification is time-independent or time dependent! This result cannot be derived from either the SEC or the DEC alone: we have exhibited a dS compactification that is compatible with the SEC but violates the DEC, and if we insist only on the DEC then it is easy to find SEC-violating matter that will allow dS-compactification (an example is a positive cosmological constant, which allows a toroidal dS compactification).

The modern understanding of Einstein’s gravitational field equations is that they arise as an effective description, valid at sufficiently low energy, of some quantum gravity theory, possibly String/M-theory. In this context, singularities indicate a breakdown of this low-energy effective description. The GMN no-go theorem is therefore a statement about limitations of the effective theory on the assumption that all matter satisfies the SEC: no non-singular time-independent dS compactification metric will solve the Einstein field equations. The no-go theorem proved here, based on the DEC, makes a subtly different statement: any non-singular time-dependent dS compactification metric that solves the Einstein field equations within some initial time period will have a singularity in its future.

The broad-brush implication of the DEC for time-dependent dS compactifications is not that it is impossible but that it can only be properly considered in the context of some ultra-violet completion of higher-dimensional General Relativity, such as String/M-theory.

²This restriction was not stated as a premise but is essential to the proof, as is an “unaveraged” implementation of the Einstein-frame condition.

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